Multifrequency study of a new Fe-rich supernova remnant in the Large Magellanic Cloud, MCSNR J0508−6902

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ABSTRACT

We present a detailed radio, X-ray and optical study of a newly discovered Large Magellanic Cloud (LMC) supernova remnant (SNR) which we denote as MCSNR J0508−6902. Observations from the Australia Telescope Compact Array (ATCA) and the XMM–Newton X-ray observatory are complemented by deep Hα images and Anglo-Australian Telescope AAOmega spectroscopic data to study the SNR shell and its shock ionization. Archival data at other wavelengths are also examined. The remnant follows a filled-in shell-type morphology in the radio continuum and has a size of $\sim 74$ pc $\times$ $57$ pc at the LMC distance. The X-ray emission exhibits a faint soft shell morphology with Fe-rich gas in its interior – indicative of a Type Ia origin. The remnant appears to be mostly dissipated at higher radio-continuum frequencies leaving only the south-eastern limb fully detectable while in the optical it is the western side of the SNR shell that is clearly detected. The best-fitting temperature to the shell X-ray emission ($kT = 0.41^{+0.05}_{-0.06}$ keV) is consistent with other large LMC SNRs. We determined an O/Fe ratio of $<21$ and an Fe mass of $0.5–1.8$ $M_\odot$ in the interior of the remnant, both of which are consistent with the Type Ia scenario. We find an equipartition magnetic field for the remnant of $\sim 28$ $\mu$G, a value typical of older SNRs and consistent with other analyses which also infer an older remnant.

Key words: ISM: supernova remnants – Magellanic Clouds – radio continuum: ISM – X-rays: ISM.

1 INTRODUCTION

The study of supernova remnants (SNRs) is an imperative part in our understanding of the Universe, as they have a profound effect on their surrounding environment. SNRs distribute heavy elements throughout the galaxy, with thermonuclear (Type Ia) explosions contributing a significant amount of iron into the surrounding medium, whilst core-collapse (CC) supernovae (SNe) inject large amounts of oxygen and other $\alpha$ elements (Matteucci & Greggio 1986). Type Ia SNe are thought to arise when a C–O white dwarf (in a binary system) approaches the Chandrasekhar limit. This occurs either through accretion from the binary companion or by a merger with a white dwarf companion, and is currently debated (see

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e.g. Maoz 2008 and references therein). CC SNe are believed to be the result of a progenitor star (with a main-sequence mass above 8 $M_{\odot}$) being drained of its nuclear fuel and therefore not being able to hold up against its own gravity through the release of nuclear energy.

The Large Magellanic Cloud (LMC) is an irregular dwarf galaxy located in close proximity (50 kpc; di Benedetto 2008) to our own Galaxy, allowing modern instruments to reach high sensitivity levels with sufficient angular resolution, from radio to X-ray wavelengths. At this distance, and because of its moderate inclination angle of 35° (van der Marel & Cioni 2001), we are able to assume that objects which lie within it are all at a similar distance. This allows us to analyse the surface brightness and spatial extent of an object accurately, which is often difficult in our own Galaxy due to uncertainties in distance, foreground absorption and background confusion. Furthermore, the LMC is considered an almost ideal environment for the study of a wide range of celestial objects due to the presence of active star-forming regions and its location outside of the Galactic plane, where absorption by gas and dust is reasonably low.

Several multiwavelength SNR selection criteria exist in the literature (e.g. Fesen, Blair & Kirshner 1985) as well as the powerful, newly evaluated emission-line diagnostic criteria reported by Frew & Parker (2010) and updated in Sabin et al. (2013). These are based on the well-known Sabbadin, Minello & Bianchini (1977) $H\alpha/[N II]$ versus $H\alpha/[S II]$ diagnostic diagram and the Baldwin, Phillips & Terlevich (1981) $[O III]5007/H\beta$ versus $[S II]/H\alpha$ diagnostic diagrams. In this paper, we adopt the criteria used by the Magellanic Cloud Supernova Remnant (MCSNR) Database which are effectively compatible with these publications. In this classification system, an SNR must satisfy at least two of the following three observational criteria: significant $H\alpha$, $[S II]$ and/or $[O III]$ optical line emission with an observed $[S II]/H\alpha$ flux ratio of $>0.4$ (Mathewson & Clarke 1973), extended non-thermal radio emission indicative of synchrotron radiation (Filipovic et al. 1998) and extended thermal X-ray emission.

The ROSAT position sensitive proportional counter (PSPC) catalogue of X-ray sources in the LMC (Haberl & Pietsch 1999) contains the object [HP99] 791. While this source was confirmed as extended, its large off-axis angle prevented a reliable classification at that time. In this paper, we present new radio-continuum and X-ray observations of [HP99] 791, from the Australia Telescope Compact Array (ATCA) and XMM–Newton, respectively. We also present for the first time the clear identification of an unambiguous complete optical shell for the remnant from a deep arcsecond-resolution $H\alpha$ image, a multi-exposure stack of the central 25 square degrees of the LMC (Reid & Parker 2006a). The shell can also be seen at lower resolution in the Magellanic Cloud Emission Line Survey (MCELS) data. Finally, we present new optical spectra from several components on the enhanced western side of the optical shell taken by the AAOmega spectrograph on the Anglo-Australian Telescope (AAT; Saunders et al. 2004; Sharp et al. 2006) which confirm a strong signature of shock excitation as expected for an SNR. This new discovery clearly satisfies all three criteria for identification as a bona fide SNR and we assign to the source the identifier MCSNR J0508−6902 following the MCSNR Database nomenclature.

The observations used in this project are described in Section 2, followed by the data analysis in Section 3. The results from these observations are explained in Section 4 and then discussed in Section 5.

### 2 Observations

#### 2.1 Radio

The ATCA is a radio interferometer located in northern New South Wales, Australia. Its new 2 GHz bandwidth and improved sensitivity have made it possible to obtain observations of a previously poorly studied group of SNRs, the older and more evolved SNRs that were not detected so far. We observed MCSNR J0508−6902 on 2011 November 15 and 16 with the ATCA, using the new compact array broadband backend receiver in the array configuration EW367, at wavelengths of 3 and 6 cm ($\nu = 9000$ and 5500 MHz). Baselines formed with the sixth ATCA antenna were omitted, as the other five antennas were arranged in a compact configuration. The observations were carried out in the so-called snap-shot mode, totalling ~50 min of integration over a 14 h period. The source PKS B1934−638 was used for primary (flux) calibration and the source PKS B0530−727 was used for secondary (phase) calibration. The phase calibrator was observed for a period of 2 min every half hour during the observations.

In addition to our own observations, we made use of a 36 cm (843 MHz) Molonglo Observatory Synthesis Telescope (MOST) mosaic image (as described in Mills et al. 1984), a 20 cm (1400 MHz) mosaic image (Hughes et al. 2007) as well as 6 cm (4800 MHz) and 3 cm (8640 MHz) mosaic images from project C918 (Dickel et al. 2005).

#### 2.2 X-ray

Haberl & Pietsch (1999) detected this source in their compilation of ROSAT PSPC observations, recording an X-ray position of RA(J2000) = 05° 08′ 36″, Dec.(J2000) = −69° 02′ 54″, and an extent of 43″ 6. It is listed as [HP99] 791. However, due to the large off-axis angle of the source, a classification was not possible. We observed MCSNR J0508−6902 with XMM–Newton on two occasions: in 2010 May (Obs. ID 0651880201, PI: M. Sasaki), and, more recently, in 2012 September (Obs. ID 0690752001, PI: P. Haberl) as part of the LMC survey. The European Photon Imaging Camera (EPIC) on board XMM–Newton (Jansen et al. 2001) consists of a pn CCD (Strüder et al. 2001) and two metal oxide semiconductor (MOS) CCD imaging spectrometers (Turner et al. 2001). The observational data were reduced using the standard reduction tasks of SAS 2.0.1, filtering for periods of high particle background. Unfortunately, Obs. ID 0651880201 was blighted by very high background, resulting in only ~9 ks of the exposure being useful. However, in the case of Obs. ID 0690752001, there was virtually no background contamination with almost all data (~25 ks) available for further analysis.

#### 2.3 Optical: Imaging and spectroscopy

We have used deep, arcsecond-resolution, narrow-band optical $H\alpha$ images created from a median stack of a dozen 2 h exposures on the central 25 square degrees of the LMC as described in detail by Reid & Parker (2006a). These data have been used as the basis for unprecedented new discoveries of emission-line sources in general and planetary nebulae (PN) in particular across the central LMC (e.g. Reid & Parker 2006a,b, 2012). The survey filter, detector and other parameters are described by Parker et al. (2005) but in essence the LMC multiple exposures were taken with fine-grained Kodak Tech-Pan films on the UK Schmidt Telescope using

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1 See http://www.mcsnrr/about.aspx. Database maintained by R. Williams et al.

2 Science Analysis Software, see http://xmm.esac.esa.int/sas/

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Downloaded from http://mnras.oxfordjournals.org/ by guest on September 10, 2014
a monolithic interference filter of exceptional quality (e.g. Parker & Bland-Hawthorn 1998). These exposures were scanned with the SuperCOSMOS measuring machine (Hambly, Irwin & MacGillivray 2001), [included with an in-built accurate world coordinate system (WCS)] transformed into electronic form and then median stacked to reach depths of $R_{\text{eqav}} \sim 22$ for Hα ($4.5 \times 10^{-17}$ erg cm$^{-2}$ s$^{-1}$ Å$^{-1}$). We extracted Hα and matching SR (short ‘broad-band’ red) fit images over a 6 arcmin × 6 arcmin area centred on the SNR with 10 µm (0.67 arcsec) pixels from the median-stacked images to create a quotient image (Hα divided by the SR), as shown in Fig. 1, reproduced to clip to the inner 98 per cent of intensity values. North-east is to the top left. A complete oval shell can be seen in Hα but with the western side being enhanced. The centroid of this oval shell and hence the best estimate for the optical centre of the remnant is RA(J2000) = 05h 08m 33s 7 and Dec.(J2000) = −69º 02´ 33˝. The major axis is 4.5 arcmin in extent and the minor axis only 4 arcmin. The positions of the two-degree field (2D) spectral samples obtained are also indicated in the figure (see below for further details).

We also checked for detection of the SNR in the MCELS that was carried out with the 0.6 m University of Michigan/Cerro Tololo Inter-American Observatory (CTIO) Curtis Schmidt telescope, equipped with a SITE 2048 × 2048 CCD, which gave a field of 1:35 at a scale of 2.3 arcsec pixel$^{-1}$ (∼5 arcsec resolution). Both the LMC and Small Magellanic Cloud (SMC) were mapped in narrow bands corresponding to Hα, [O III] (λ = 5007 Å) and [S II] (λ = 6716, 6731 Å). All the data have been flux calibrated and assembled into mosaic images, a small section of which is shown in Fig. 2. Further details regarding the MCELS are given by Smith, Points & Winkler (2006) and at http://www.ctio.noao.edu/mcels. The MCELS data also reveal the SNR though at the poorer angular resolution it is difficult to discern the full oval optical shell. However, MCELS does have the added advantage of the [O III] and [S II] filters providing additional diagnostic capabilities. The composite RGB image with (R = Hα, G = [S II] and B = [O III]) centred on the new SNR is shown in Fig. 2 with the XMM–Newton contours in red (0.7–1.1 keV) and ATCA 5500 MHz radio data contours in white. Interestingly, the strongest radio signature corresponds to the weakest optical emission to the south-east while the X-ray emission occupies the centre of the shell.

### 2.3.1 AAT AAOmega spectroscopy

As part of an AAOmega 2dF service night on the AAT on 2013 May 1, 2 × 900 s exposures were taken on an LMC 2dF centred at RA(J2000) = 05h 13m 15s 7 and Dec.(J2000) = −69º 29´ 51˝ for the LMC PN project of Reid and Parker (e.g. Reid & Parker 2006a,b, 2013). AAOmega is a powerful double-beam spectrograph (Sharp et al. 2006) that is coupled to the 2dF multi-object fibre spectroscopy system on the AAT (Lewis et al. 2002). Due to the 400 available 2dF fibres, it was possible to sacrifice five 2 arcsec diameter fibres to allocate to the new SNR candidate that happened to be in the field of view (FOV) of one of the Reid–Parker LMC PN fields. The individual SNR fibres were placed at carefully located positions on the western and eastern edges of the SNR shell as indicated in Fig. 1 with their accurate positions listed in Table 1.

Both the 580V and 1000R gratings were simultaneously employed for the general LMC PN spectroscopic observations with the 1000R red grating centred at 6850 Å and giving coverage from 6300 to 7425 Å. This is ideal for medium-resolution coverage of the diagnostic Hα, [N II] and [S II] nebular emission lines common to both PN and SNRs and also provides for more accurate radial velocity estimates. The lower resolution 580V blue grating was centred at 4800 Å giving coverage from 3700 to 5850 Å. The blue data suffered from the effects of extinction and low signal-to-noise ratio (S/N) and were of limited use for the SNR observations.

The AAOmega data were reduced with the standard 2df data reduction (2DFDR) pipeline (e.g. Croom, Saunders & Heald 2004). Note that the optical fibres sample both target and superposed sky. Hence, it is standard practice to allocate a certain fraction of the 400 2dF fibres to blank sky. This is normally ∼20 fibres as was
3 DATA ANALYSIS

3.1 Radio

The MIRIAD (Sault, Teuben & Wright 1995) and KARMA (Gooch 1995) software packages were used for reduction and analysis. More information on the observing procedure and other sources observed in this session/project can be found in Bouzzeto et al. (2012a,b,c, 2013) and de Horta et al. (2012).

Images were formed using MIRIAD multifrequency synthesis (Sault & Wieringa 1994) and natural weighting. They were deconvolved using the MFCLEAN and RESTOR algorithms with primary beam correction applied using the limnos task. The 6 cm (5500 MHz) ATCA image together with 20 cm (1337 MHz) and 36 cm (843 MHz) mosaic images of MCSNR J0508−6902 is shown in Fig. 4. The resolution of the 6 cm image is 37.7 arcsec × 25.1 arcsec.

3.2 X-rays

3.2.1 Imaging

We extracted images and exposure maps in three energy bands suited to the analysis of SNRs, whose X-ray emission is mostly thermal: a soft band (0.3−0.7 keV) including strong lines from oxygen; a medium band (0.7−1.1 keV), which includes Fe L-shell lines, as well as the Heα and Lyα lines of Neon; and a hard band (1.1−4.2 keV) comprising lines from Mg, Si, S, Ca, Ar and, if present, non-thermal continuum.

From the pn detector, we selected single- and double-pixel events (PATTERN = 0 to 4). Below 0.5 keV, only single-pixel events were selected to avoid the higher detector noise contribution from the double-pixel events. All single- to quadruple-pixel (PATTERN = 0 to 12) events from the MOS detectors were used. Bad pixels were masked. The detector background (see Section 3.2.3) was taken from filter wheel closed (FWC) data.3 Before it was subtracted from the raw images, we scaled it by a factor estimated from the count rates in the corners of the detectors not exposed to the sky.

We merged MOS and pn data, and combined images from the two XMM–Newton observations. We then adaptively smoothed the images, computing Gaussian kernels at each position in order to reach an approximate S/N of 5. The minimum size for the kernels was set to an FWHM of 20 arcsec. Smoothed images were finally divided by the vignette exposure maps.

Combining the final images, we show in Fig. 5 the XMM–Newton ‘true-colour’ image of MCSNR J0508−6902, which reveals a striking morphology. A bright central region, dominated by X-rays in the medium band, is surrounded by a fainter and softer shell. The brightness of the interior emission in the 0.7−1.1 keV band, which includes the Fe L-shell lines, suggests that MCSNR J0508−6902 falls into the class of SNRs with central Fe enhancements.

Examples of these objects include 0548−70.4 and 0534−69.9 (Hendrick, Borkowski & Reynolds 2003), DEM L238 and DEM L249 (Borkowski, Hendrick & Reynolds 2006) and SNR 0104−72.3 (Lee et al. 2011; Roper et al., in preparation). Spectral differences between the central region and the shell are analysed in Section 4.3. There is little emission from the object above 1.1 keV. Several point sources are
Figure 3. 1000R red grating 1D 2dF spectra for fibres 238, 248 and 260 taken on the enhanced western edge and a combined 1D spectrum from fibres 253 and 259 taken on the lower surface brightness eastern edge as indicated in Fig. 1.

Table 2. Heliocentric radial velocities estimated from the reduced 1D 1000R red individual AAOmega fibre spectra as provided by the IRAF emsao package.

<table>
<thead>
<tr>
<th>2dF fibre ID</th>
<th>Radial velocity (km s(^{-1}))</th>
<th>Error (km s(^{-1}))</th>
<th>Comments</th>
</tr>
</thead>
<tbody>
<tr>
<td>Fib 238</td>
<td>287.8</td>
<td>6.4</td>
<td>for 2/5 lines</td>
</tr>
<tr>
<td>Fib 248</td>
<td>323.0</td>
<td>14.4</td>
<td>for 5/5 lines</td>
</tr>
<tr>
<td>Fib 260</td>
<td>275.0</td>
<td>15.1</td>
<td>for 3/5 lines</td>
</tr>
<tr>
<td>Fib 253</td>
<td>263.5</td>
<td>13.7</td>
<td>for 3/6 lines</td>
</tr>
<tr>
<td>Fib 259</td>
<td>267.8</td>
<td>21.3</td>
<td>for 5/6 lines</td>
</tr>
</tbody>
</table>

Figure 4. 843 MHz (top), 1400 MHz (middle) and 5500 MHz (bottom) radio-continuum images of MCSNR J0508−6902 with an overlaid ellipse indicating the approximate optical association for this SNR. The side bar quantifies the total intensity in mJy per beam.)
The contours are from 3σ Gaussian width of 20 arcsec, and overlaid with 1400 MHz radio contours.

0.7–1.1 keV, blue: 1.1–4.2 keV) of MCSNR J0508−6902 smoothed with a Gaussian width of 20 arcsec, and overlaid with 1400 MHz radio contours. The contours are from 3σ to 6.5σ in steps of 0.5σ.

detected outside the boundary of the remnant, which are most likely background objects behind the LMC.

3.2.2 X-ray spectral analysis

We consider only Obs. ID 0690752001 for our spectral analysis. The ∼9 ks of data from Obs. ID 0651880201 can add little to our analysis, as e.g. the faint shell of the remnant is not detected in images from this observation only. Furthermore, using the diagnostic devised by De Luca, Molendi, & Leccardi (2004), we found that the flare-filtered event lists were still heavily contaminated by soft protons (SPs).

In addition, only EPIC-pn spectra are considered as the SNR shell is extremely faint in the EPIC-MOS images. Before proceeding with the spectral extraction, we generated a vignetting-weighted event list to correct for the effective area variation across our extended source using the SAS task evigweight. Only single- and double-pixel events (PATTERN = 0 to 4) are considered. The extraction area of the central region was defined by enclosing the 0.7–1.1 keV band contours (see Fig. 2) with an ellipse (~2.8 arcmin × 1.6 arcmin). The shell spectrum was extracted from an ~5 arcmin diameter region (as determined in Section 4.2), excluding the central extraction region. Backgrounds should ideally be chosen from nearby positions on the detector. However, since the SNR was located close to on-axis and nearby suitable regions were comparatively source- and chip-gap heavy, this was not possible. As such, backgrounds were extracted from a source and diffuse emission-free region immediately to the north-west of the remnant, further off-axis than the source. This introduces important effects which must be considered in the spectral modelling of the particle-induced background, which is discussed in the forthcoming text. All spectra were rebinned using the ftool grppha so that each bin contained a minimum of 30 counts. All fits were performed using XSPEC (Arnaud 1996) version 12.7.1 with abundance tables set to those of Wilms, Allen & McCray (2000), and photoelectric absorption cross-sections set to those of Balucinska-Church & McCammon (1992).

3.2.3 X-ray background

Spectral analysis of extended X-ray sources is complicated by the contributions of the X-ray background, comprising particle-induced and astrophysical components. Often, a subtraction of a nearby background region is inappropriate due to the spectral and/or spatial variation of these background components. While the consideration of the X-ray background for bright extended sources may not be so significant, in the case of faint emission (as is the case with the shell of MCSNR J0508−6902 in particular), the background must be carefully accounted for in order to extract the purest possible information from the source. A commonly used method (see Maggi et al. 2012 for a recent example) is to extract a nearby background spectrum, define a physical model for the background and simultaneously fit this model to the source and background spectra. We adopt this method for our analysis of MCSNR J0508−6902. A detailed description of our background treatment is given in Appendix A.

3.2.4 MCSNR J0508−6902 emission

Line-of-sight absorption to MCSNR J0508−6902 consists of a Galactic component (N_HGal, fixed at 7 × 10^{20} cm^{-2} as described in Appendix A3) and a foreground LMC component (N_HLMC). These were modelled using photoelectric absorption components with Galactic and LMC abundances, respectively.

Shell emission. We modelled the shell emission using various thermal plasma models appropriate for SNRs. We report here on the results using the standard Sedov model (Borkowski, Lyerly & Reynolds 2001), implemented in XSPEC as vsedov. Given the large size of the remnant, the assumption that it is in the Sedov phase is justified. The free parameters of the vsedov model are the mean shock temperature (kT_s), post-shock electron temperature (kT_e), metal abundances, ionization time-scale (τ_0, the electron density immediately behind the shock front multiplied by the age of the remnant). Due to the relatively poor count statistics, metal abundances for the model were fixed at the average LMC values (0.5 solar; Russell & Dopita 1992). In the case of MCSNR J0508−6902, fixing the metal abundances at 0.5 solar is quite reasonable as we expect the shell to be dominated by swept-up interstellar medium (ISM).

Interior emission. Fits of the central spectrum are complicated by the fact that we must consider both the interior Fe-rich gas and the emission from the shell overlapping this region. To account for the shell emission, we included a vsedov component in the fits to the central spectrum with the shell kT_s (= kT_e) and τ_0 fixed to the best-fitting values determined from fits to the shell spectrum (see above). For the interior emission itself, we considered several thermal plasma models, which are discussed in detail in Section 4.3.2.

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See [http://xmm2.esac.esa.int/external/xmm_sw_cal/background/epic_scripts.shtml](http://xmm2.esac.esa.int/external/xmm_sw_cal/background/epic_scripts.shtml)

See [http://heasarc.nasa.gov/lheasoft/ftools/fhelp/grppha.txt](http://heasarc.nasa.gov/lheasoft/ftools/fhelp/grppha.txt)
measurements (and potentially closer to what we see in the zero-spacing measurements). We use the spectral index which is inclusive of zero-spacing but loses the shell emission of the remnant until we get to 9000 MHz, at which point all emission from the SNR is lost.

The large extent of MCSNR J0508−6902 made a spectral energy distribution (SED) estimate difficult, as observations using an interferometer suffer from missing zero-spacings (which are responsible for large-scale, extended emission), and therefore, missing flux. Because of this, we take two different sets of flux density measurements to estimate the SED. The first set of measurements came from mosaic images (20 cm from Hughes et al. 2007 and 6 cm from Dickel et al. 2010) which include a zero-spacing observation. The second set (which do not use a zero-spacing observation) include the two aforementioned mosaics before the zero-spacing (Parkes) observations were added, in addition to a 36 cm measurement from the MOST mosaic image (as described in Mills et al. 1984). The values of these flux density measurements are reported in Table 3.

Table 3. Integrated flux densities of MCSNR J0508−6902 measured in this study.

<table>
<thead>
<tr>
<th>$\lambda$ (cm)</th>
<th>$\nu$ (MHz)</th>
<th>rms (mJy)</th>
<th>Beam size (arcsec)</th>
<th>$S_{\text{Total}}$ (mJy)</th>
<th>$\Delta S_{\text{Total}}$ (mJy)</th>
</tr>
</thead>
<tbody>
<tr>
<td>20</td>
<td>1377</td>
<td>0.7</td>
<td>40.0 × 40.0</td>
<td>88</td>
<td>15</td>
</tr>
<tr>
<td>6</td>
<td>4800</td>
<td>1.0</td>
<td>35.0 × 35.0</td>
<td>41</td>
<td>10</td>
</tr>
<tr>
<td>36</td>
<td>843</td>
<td>0.4</td>
<td>46.4 × 43.0</td>
<td>59</td>
<td>6</td>
</tr>
<tr>
<td>20</td>
<td>1377</td>
<td>0.7</td>
<td>40.0 × 40.0</td>
<td>51</td>
<td>5</td>
</tr>
<tr>
<td>6</td>
<td>4800</td>
<td>–</td>
<td>33.0 × 33.0</td>
<td>19</td>
<td>2</td>
</tr>
</tbody>
</table>

The top two flux density measurements include a zero-spacing measurement, whereas the lower three measurements do not.

4 RESULTS

4.1 Radio

This remnant exhibits a filled-in shell morphology (at 20 cm) centered at RA(J2000) = 05° 08′ 33.7" and Dec.(J2000) = −69° 02′ 33″. There are two unrelated background sources in the field, one in the south [RA(J2000) = 05° 07′ 57.7", Dec.(J2000) = −69° 03′ 59.49", and the other to the west [RA(J2000) = 05° 08′ 06′.94", Dec.(J2000) = −69° 01′ 51.12"] (Fig. 4). At lower radio frequencies in this study (843 and 1400 MHz), the emission of the SNR is fairly uniform across the remnant with significant limb brightening in the south-east. As we move to higher frequencies (≥5500 MHz), the remnant maintains this south-eastern limb but loses the shell emission of the remnant until we get to 9000 MHz, at which point all emission from the remnant is lost.

MCSNR J0508−6902 is similar to that of the 20 cm radio-continuum emission (Fig. 5), filling out the shell of the remnant. X-ray emission in the medium band (0.7–1.1 keV) appears to be encased within the radio and optical emission, making the central location of this emission evident.

In the MCELS combined RGB Hα, [O iii] and [S ii] optical image, we can see an association between the [O iii] emission and radio emission (Fig. 2) where the [O iii] emission perfectly fits within the confines of the radio emission along the south-eastern limb of the remnant and then extending in the northern region to further complete the ellipse ring structure of the remnant. On the other hand, the Hα and [S ii] emission located in the north-western regions of the remnant shows no association with the radio-continuum emission at any frequency, or with the X-ray emission (Fig. 2). The Hα and [S ii] emission in the north-west is significantly stronger when compared to the faint [O iii] emission. Strangely enough, they almost perfectly complement the [O ii] emission in creating an ellipse ring morphology. This morphology is similar to that of LMC SNR J0453−6829 (Haberl et al. 2012).

Due to the higher resolution of Hα images from the deep Hα survey of the central 25 square degrees of the LMC by Reid et al. 2010, we were able to estimate the magnetic field strength for this SNR based on the equipartition formula as given in Arbutina et al. (2012). Equipped with the measured magnetic field strength, we were unable to determine the Faraday rotation and therefore can not deduce the magnetic field strength. As an alternative however, we were able to estimate the magnetic field strength for this SNR based on the equipartition formula as given in Arbutina et al. (2012).

$\alpha = -0.69 \pm 0.12$ (Without Parkes)

$\alpha = -0.62 \pm 0.34$ (With Parkes)

Figure 6. Radio-continuum spectrum of MCSNR J0508−6902.
Parker (2006a) (Figs 1 and 7), we are able to see a complete optical shell which perfectly follows the radio emission in the south-east. The enhanced optical emission seen in this image towards the north-west suggests propagation of the shock front into the slightly higher density ISM, compared with somewhat lower density ISM on the south-east side. Separate fragmented Hα filaments inside the shell are common in optical observations (as well as in radio), as Hα emission is dominant where the blast propagates into dense, cooler regions.

X-ray, radio and optical images were combined to estimate the extent of this object. To do so, we created an annotation file which comprised of an ellipse that followed the outer emission from these combined images. We deduce an extent of 304 arcsec × 234 arcsec ± 4 arcsec (74 × 57 ± 1 pc), which places this remnant amongst the 20 per cent largest SNRs in the LMC (Badenes, Maoz & Draine 2010).

To gain an understanding of the environment in which the remnant resides, we use an image from the NANTEN molecular cloud survey by Kawamura et al. (2009), and another from the higher resolution Magellanic Mopra Assessment (MAGMA) survey by Wong et al. (2011). From these images (Fig. 8), it appears that the remnant is confined within the CO emission, elongated and stretching into the less dense region towards the north-east. However, this may be a projection effect as there is no reason to believe that the SNR is at the same distance as the molecular cloud. Indeed, the molecular cloud shown in Fig. 8 is most likely located, and may have been formed in SGS 5 (Dawson et al. 2013). The extent of SGS 5 through the LMC disc increases the likelihood of an interaction with MCSNR J0508—6902 and can provide the density gradient required to explain the asymmetric optical and radio emission.

Because of the bright interior X-ray emission, this remnant could be suggested to be a member of the mixed-morphology SNR (MMSNR; Rho & Petre 1998; Lazendic & Slane 2006). These remnants show indications for interactions with molecular clouds and often show associated OH maser emission. These features led some authors to conclude to a CC origin for these type of SNRs (e.g. Miceli 2011). In the case of MCSNR J0508—6902, although there is some evidence of interaction with a molecular cloud, we clearly favour a Type Ia SN progenitor (see Section 4.3.2). We note that other Type Ia SNRs have been found in a molecular cloud/star forming environment (e.g. Lewis et al. 2003; Lee et al. 2011). In addition, overionized (i.e. recombining) plasmas have been found recently in several Galactic MMSNRs (Yamaguchi et al. 2009; Ohnishi et al. 2011; Yamauchi et al. 2013), while this feature is absent in the spectrum of MCSNR J0508—6902. Finally, we note that this classification has only been applied to Galactic remnants and not to the (usually Fe-rich) ‘centrally peaked in X-rays’ SNRs in the Magellanic Clouds (see Section 5 and references therein).

4.3 X-ray emission of MCSNR J0508—6902

4.3.1 Shell emission

As stated in Section 3.2.4, we modelled N_H_LMC using a photoelectric absorption model with LMC abundances. However, N_H_LMC consistently tended to zero, with upper 90 per cent confidence...
intervals for this parameter of $\sim 8 \times 10^{20} \text{cm}^{-2}$ for all fits. Thus, we fixed $N_{\text{H},\text{LMC}}$ at 0. Initial shell fits with the vsedov model resulted in best-fitting values for $kT_s$ and $kT_e$ to be equal within their 90 per cent confidence intervals and, thus, $kT_s$ and $kT_e$ were constrained to be the same. Indeed, we should expect that evolved SNRs are close to ion–electron temperature equilibrium (Borkowski et al. 2001).

The results of the spectral fits to the shell and central regions of MCSNR J0508−6902 are given in Table 4, with the spectra and best-fit models plotted in Fig. 10. The vsedov model provides an acceptable fit to the shell spectrum with a reduced $\chi^2$ of 1.21. The best-fitting temperature to the shell emission of $kT_s = kT_e = 0.41(0.35–0.46)$ keV is consistent with other large LMC SNRs (see Williams et al. 2004; Grondin et al. 2012, for example), as well as DEM L238 and DEM L249 from Borkowski et al. (2006).

The high value of $\tau_0$ ($\sim 10^{12}$ s cm$^{-3}$) indicates that the plasma is in ionization equilibrium. From the shell component in the fits to the central spectra, we determined that 13–28 per cent of the total shell emission is contributing to the central spectrum. The absorption-corrected X-ray luminosity ($L_X$) of the shell in the 0.2–10 keV range was found to be $\sim 5 \times 10^{34}$ erg s$^{-1}$.

Using our fit results we can further estimate physical parameters of the SNR using the Sedov dynamical model (see Sasaki et al. 2004, for example). We know that MCSNR J0508−6902 is not spherically symmetric since the semi-major and semi-minor axes are 37 pc and 28.5 pc, respectively. Taking these semi-major and semi-minor axes as the first and second semi-principal axes of an ellipsoid describing MCSNR J0508−6902, and assuming that the third semi-principal axis is in the range 28.5–37 pc, we determined the volume ($V$) limits for the remnant and their corresponding effective radii ($R_{\text{eff}}$) to be $(3.7–4.8) \times 10^{60}$ cm$^3$ and 31.1–33.9 pc, respectively. In the following calculations, we evaluate the physical parameters of the remnant at these volume and effective radius limits, and determine the resulting range for each parameter.

The X-ray temperature corresponds to a shock velocity

$$v_s = \left( \frac{16 kT_s}{3 \mu} \right)^{0.5},$$

where $\mu$ is the mean molecular weight of the gas.

**Table 4.** Results of the spectral fits. For the vsedov model, the post-shock mean temperature and post-shock electron temperature were set to be the same ($kT_s = kT_e = kT$) and the abundance is set to that of the LMC (0.5 solar). The EM and $L_{X,0.2–10 \text{ keV}}$ values for the shell were calculated for the total shell (i.e. central and shell region). The ranges in parentheses represent the 90 per cent confidence intervals on the parameters.

<table>
<thead>
<tr>
<th>Fit parameter/result</th>
<th>Value</th>
</tr>
</thead>
<tbody>
<tr>
<td>Shell emission: vsedov model</td>
<td></td>
</tr>
<tr>
<td>$kT$ (keV)</td>
<td>0.41 (0.35–0.46)</td>
</tr>
<tr>
<td>$\tau$ (10$^{12}$ s cm$^{-3}$)</td>
<td>2.70 (1.73–6.92)</td>
</tr>
<tr>
<td>$\chi^2/\nu$</td>
<td>267.08/219</td>
</tr>
<tr>
<td>EM (10$^{57}$ cm$^{-3}$)</td>
<td>4.75 (2.54–7.72)</td>
</tr>
<tr>
<td>$L_{X,0.2–10 \text{ keV}}$ (10$^{34}$ erg s$^{-1}$)</td>
<td>5.1</td>
</tr>
<tr>
<td>Interior emission: pure heavy-element vapec models</td>
<td></td>
</tr>
<tr>
<td>$kT_{Fe,0}$ (keV)</td>
<td>0.78 (0.75–0.81)</td>
</tr>
<tr>
<td>$\chi^2/\nu$</td>
<td>66.49/64</td>
</tr>
<tr>
<td>$n_{Fe,0}$ (10$^{53}$ cm$^{-3}$)</td>
<td>0.58 (0.54–0.61)</td>
</tr>
<tr>
<td>$n_{O,0}$ (10$^{53}$ cm$^{-3}$)</td>
<td>2.99 (0–9.96)</td>
</tr>
<tr>
<td>$L_{X,0.2–10 \text{ keV}}$ (10$^{34}$ erg s$^{-1}$)</td>
<td>2.8</td>
</tr>
</tbody>
</table>
where $kT_s$ is the post-shock temperature and $\mu$ is the mean mass per free particle. For a fully ionized plasma with LMC abundances, $\mu = 0.61$, resulting in a shock velocity of $v_s = 572(\pm 36)$ km s$^{-1}$. The age of the remnant can now be determined from the similarity solution:

$$v_s = \frac{2R}{5t},$$

where $R = R_{\text{eff}}$ and $t$ is the age of the remnant. This gives an age range of $20-25$ kyr. The pre-shock H density ($n_{\text{H}1,0}$) in front of the blast wave can be determined from the emission measure (EM, see Table 4). Evaluating the emission integral over the Sedov solution using the approximation for the radial density distribution of Kahn (1975) gives

$$\int n_e n_{\text{H}1} dV = EM = 2.07 \left( \frac{n_0}{n_{\text{H}1,0}} \right) n_{\text{H}1,0} V,$$

where $n_e$ and $n_{\text{H}1}$ are electron and hydrogen densities, respectively, and $V$ is the volume (e.g. Hamilton, Sarazin & Chevalier 1983). Taking $n_0/n_{\text{H}1,0} = 1.21$ and the determined range of volumes, this equation yields $n_{\text{H}1,0} = (1.5-2.8) \times 10^{-2}$ cm$^{-3}$. Since the pre-shock density of nuclei is given as $n_0 \sim 1.1 n_{\text{H}1,0}$, it follows that $n_0 = (1.7-3.1) \times 10^{-2}$ cm$^{-3}$. Hence, the SNR is expanding into a fairly rarefied environment. The initial explosion energy ($E_0$) can be determined from the equation:

$$R = \left( \frac{2.02 E_0 \rho_0^2}{\mu m_p} \right)^{1/5},$$

where $\mu_0$ is the mean mass per nucleus (=1.4$m_p$). This results in an initial explosion energy in the range $(0.37-0.75) \times 10^{51}$ erg which is less than the canonical value of $10^{51}$ erg but consistent with several Type Ia SNRs in the LMC (see Hendrick et al. 2003, for example). The swept-up mass contained in the shell is given simply by $M = V \mu_0 n_0$, which is evaluated to be $(32-62) M_{\odot}$. This large amount of swept-up material supports our original assumption that MCSNR J0508–6902 is into the Sedov phase of its evolution. All the derived parameters of MCSNR J0508–6902 are summarized in Table 5. Using the determined values of $t$ and $n_0$, we estimated the expected ionization age for the remnant, given as $4.8 n_{\text{H}1,0} t$, to be in the range $(5.3-9.2) \times 10^6$ s. This is significantly less than the ionization age determined from the spectral fits with the Sedov model. Given this inconsistency, it is likely that the Sedov model may not be strictly applicable to MCSNR J0508–6902. Several other remnants exhibit a similar discrepancy between measured and expected ionization time-scales. Hughes, Hayashi & Koyama (1998) and van der Heyden et al. (2003) observed this for candidate CC SNRs in the LMC and SMC, respectively. Hughes et al. (1998) suggest the explanation to be that the remnants initially expand into a low-density cavity in the ISM, likely due to the massive stellar wind. The remnant rapidly reaches the wall of the cavity where it meets the denser material, slows down and emits X-rays. This scenario represents a deviation from the standard Sedov model of a remnant expanding into a uniform ambient medium and can account for the discrepancy in the measured and expected ionization ages. The problem with this scenario as an explanation for MCSNR J0508–6902 is that a high-mass stellar wind is required to alter the ambient ISM in such a manner which is likely not the case for MCSNR J0508–6902. Borkowski et al. (2006) also observe such a discrepancy for DEM L238, a Type Ia remnant. They suggest that their derived ambient ISM density may be underestimated resulting in an increased ionization age in the west as compared to the east. This may explain the observed ionization parameter discrepancy. However, in such a picture we would expect variation of the X-ray emission in this direction. Unfortunately, given the already

![Figure 10.](#) XMM–Newton spectral fits to MCSNR J0508–6902 shell (left) and central spectra (right) and associated backgrounds. The shell spectrum is shown in green, the shell background in blue, the central spectrum in black and the central background in red, with the best-fitting model to each represented by the solid lines of matching colour. The astrophysical background components are indicated by the dark grey dotted lines in each plot. The particle-induced background components are indicated by the light grey dotted lines. The green dash–dotted line in each spectrum represents emission from the shell. The black and magenta dash–dotted lines in the central spectrum show the pure Fe component and the pure O component.

Table 5. Physical properties of MCSNR J0508–6902 derived from the Sedov model.

<table>
<thead>
<tr>
<th>$n_0$ (10$^{-2}$ cm$^{-3}$)</th>
<th>$v_s$ (km s$^{-1}$)</th>
<th>$t$ (kyr)</th>
<th>$M$ (M$_{\odot}$)</th>
<th>$E_0$ (10$^{51}$ erg)</th>
</tr>
</thead>
<tbody>
<tr>
<td>1.7–3.1</td>
<td>536–608</td>
<td>20–25</td>
<td>32–62</td>
<td>0.37–0.75</td>
</tr>
</tbody>
</table>
low count statistics, an analysis of spatio-spectral variation is not really feasible. In any case, because of the discrepancy between the measured and expected ionization parameters, the SNR properties derived from the Sedov dynamical model are somewhat uncertain.

4.3.2 Interior emission

We initially fitted the interior emission with a *vapec* model at LMC abundances. However, this resulted in a very poor fit (red, $\chi^2 > 2$) with large residuals at $\sim 0.9$ keV due to the Fe L-shell emission lines. Thus, we freed the Fe abundance and re-performed the fit, resulting in a substantially improved fit with $kT = 0.76(0.73-0.79)$ keV, $Z_{\text{Fe}} > 2.01 Z/Z_\odot$ and red. $\chi^2 = 1.08$. The lower limit for the Fe abundance is approximately four times higher than the average LMC value, clearly indicating an overabundance of Fe in the central region, most likely due to Fe-rich ejecta being heated by the reverse shock. Allowing other metal abundances to vary (e.g. O) did not improve the fits and were very poorly constrained.

A more realistic model can be applied assuming that the interior gas consists of ejecta only. Hence, we applied thermal plasma models, representative of pure heavy-element plasmas (see Kosenko, Helder & Vink 2010, for example). We performed trial fits with collisional ionization equilibrium (CIE) models and non-equilibrium ionization (NEI) models representing the pure metal plasmas and found that, in both cases, an Fe plasma was required to fit the spectrum of the interior emission with a minimal pure O contribution. The CIE models provided acceptable fits to the spectra (reduced $\chi^2 \approx 1$). Similar quality fits were achieved with NEI models with the formal best-fitting ionization time-scales determined as $\sim 10^{13}$ s cm$^{-3}$, suggestive of a plasma in CIE. However, no constraints could be placed on these values. Given the quality of the trial fits with the CIE models, the interior plasma is suggested to be close to or in CIE. Hence, we model this emission using the *vapec* model in *XSPEC*. Initial fits showed that the temperature of the O component was not well constrained. Hence, as in Hughes et al. (2003), we simplified the model by assuming the O and Fe to be co-spatial and constrained their temperatures to be the same. This yielded a best-fitting $kT = 0.78(0.73-0.79)$ keV. In our spectral fits, we found that Fe dominates the interior emission with O having little or no contribution to the spectrum. Indeed, the normalization of the O component was pegged at zero, demonstrating that it is formally not required. However, we keep the component in the spectral model as the upper limit to the O contribution can allow us to at least determine the maximal O/Fe ratio (see below). We reiterate that the interior plasma is likely in or close to CIE due to the quality of the *vapec* model fit (reduced $\chi^2 \approx 1$, see Table 4). The $L_X$ of the interior gas in the 0.2–10 keV range was found to be $\sim 2.8 \times 10^{34}$ erg s$^{-1}$.

It is possible to identify the likely SN mechanism responsible for MCSNR J0508–6902 by assessing the O/Fe ratio of the interior region. From the SN nucleosynthesis yields of Iwamoto et al. (1999), the O/Fe ratio is expected to be 0.3–0.7 by number for Type Ia events and $\sim 70$ for CC events. We estimated the maximum allowed O/Fe ratio for MCSNR J0508–6902 using a $\chi^2$ contour plot produced for the Fe and O normalization parameters. This yielded a maximum ratio of $< 21$ at the 90 per cent confidence level, slightly higher than the LMC value of 13.2 (Russell & Dopita 1992). We stress that this upper limit is based on the rather poorly constrained, formally not required O component in the spectral fits. However, it does rule out a CC scenario, indicating that MCSNR J0508–6902 likely resulted from a Type Ia explosion. Another diagnostic for the SN type is the mass of Fe produced. This can be estimated using equation 2 of Kosenko et al. (2010), adapted for Fe. We split this equation into terms which are easily determined from the fit results, given as

$$M_{\text{Fe}} = (\frac{V_{\text{Fe}} E M_{\text{Fe}}}{n_{\text{Fe}}} )^{0.5} (n_{\text{Fe}}/n_{\text{e}})^{0.5} m_1 A_{\text{Fe}},$$  \hspace{1cm} (5)$$

where $V_{\text{Fe}}$ is the volume occupied by the Fe, $EM_{\text{Fe}} = n_{\text{Fe}} V$ is the EM of the Fe gas, $n_{\text{Fe}}/n_{\text{e}}$ is the electron-to-Fe-ion ratio, $m_1$ is the atomic mass unit and $A_{\text{Fe}}$ is the atomic mass of Fe. $V_{\text{Fe}}$ is determined from the dimensions of the central region (1.4 arcmin × 0.8 arcmin) and assuming an ellipsoidal morphology with a third semi-axis equal to the mean of the semi-major and semi-minor axes (1.1 arcmin). If the actual morphology is oblate or prolate, the volume would be 13 per cent higher or lower, respectively, $EM_{\text{Fe}}$ is obtained from the normalization of the Fe component and is listed in Table 4. We must estimate the $n_{\text{Fe}}/n_{\text{e}}$ ratio. As discussed in Hughes et al. (2003), there are two plausible scenarios for the number of electrons per Fe ion, depending on the level of admixture of H in the ejecta. If there is no H, then the number of free electrons per Fe ion only depends on the average ionization state of the each metal contributing electrons. We do not consider the contribution of electrons by the minimal O content of the interior region. For a plasma in CIE at $kT = 0.78$ keV (log $T = 6.95$), the average ionization is 18.3 for Fe (Shull & van Steenberg 1982). Then, from equation (5), the Fe mass is in the range 1.6–1.8 $M_\odot$ for the pure heavy-element scenario. The second case proposed by Hughes et al. (2003) is that a similar mass of H is mixed into the metal-rich ejecta. Therefore, the number density of Fe over H is 1/56, and the average number of electrons per Fe ion is 74.3. Thus, from equation (5) we have an Fe mass range of 0.8–0.9 $M_\odot$ if there is a comparable mass of H to Fe in the ejecta. The former case gives a value somewhat too high for the explosion of a white dwarf, the latter value being typical of a Type Ia yield. However, the ejecta might be clumpy, with a filling factor of $\sim 0.4$ (see Kosenko et al., 2010, and references therein), in which case the Fe content reduces by a factor of $\sqrt{0.4} \approx 0.63$ giving 1.0–1.1 and 0.5–0.6 $M_\odot$ for the pure Fe and Fe+H mixture cases, respectively. The large mass of Fe present in the interior further supports a Type Ia progenitor for MCSNR J0508–6902.

5 DISCUSSION

The X-ray emission from MCSNR J0508–6902 is notable due to the Fe-rich gas in the interior of the remnant, which is typically indicative of a Type Ia origin (Vink 2012, and references therein). MCSNR J0508–6902 is not the only SNR to exhibit such a feature, with a number of SNRs also characterized by central Fe enhancements in both the LMC: 0454–67.2 (Seward et al. 2006); 0548–70.4 and 0534–69.9 (Hendrick et al. 2003); DEM L316A, DEM L238 and DEM L239 are most similar to MCSNR J0508–6902, as not only is a faint shell detected surrounding the interior Fe-rich gas, but also the Fe-rich gas is in CIE. However, as argued by Borkowski et al. (2006), it is likely that many more SNRs fall into this category, and only their shells are so faint as to be undetectable by *XMN–Newton*. DEM S128, IKT 5 (Filipović et al. 2000, 2008; van der Heyden, Bleeker & Kaastra 2004); SNR 0104–72.3 (Lee et al. 2011). Of these, DEM L316A, DEM L238 and DEM L239 are most similar to MCSNR J0508–6902, as not only is a faint shell detected surrounding the interior Fe-rich gas, but also the Fe-rich gas is in CIE. However, as argued by Borkowski et al. (2006), it is likely that many more SNRs fall into this category, and only their shells are so faint as to be undetectable by *XMN–Newton*. DEM S128, IKT 5 (Filipović et al. 2000) and IKT 5 (Roper et al., in preparation) are examples of such SNRs with an Fe-rich gas in CIE but no detected shell. The X-ray structure of such SNRs presents a problem for standard Type Ia models (Dwarkadas & Chevalier 1998; Badenes et al. 2003, 2006; Badenes, Borkowski & Bravo 2005). Assuming that the SN ejecta have...
expanded in a uniform ambient ISM, and we know from the X-ray analysis that the ambient density is quite low ($\sim 10^{-2} \text{ cm}^{-3}$), then, under the standard Type Ia models, the central Fe-rich gas should be of low density and the ionization times are short, which are not observed. Borkowski et al. (2006) demonstrate these scenarios using 1D hydrodynamic simulations using the model of Dwarkadas & Chevalier (1998) and find consistency with the results of Badenes et al. (2003, 2005, 2006).

To explain the X-ray properties of DEM L238 and DEM L249, Borkowski et al. (2006) argue that, rather than the SNR expanding into an ideal environment of uniform density, the structure of the ejecta and/or the uniformity of the ambient ISM deviates from the standard Type Ia models. The best candidates for such a scenario are the ‘prompt’ Type Ia explosions (Mannucci, Della Valle & Panagia 2006; Sullivan et al. 2006; Aubourg et al. 2008). These ‘prompt’ events are produced by relatively massive stars ($\sim 3.5$–$8 \text{ M}_\odot$; Aubourg et al. 2008) and are capable of increasing the circumstellar medium (CSM) density via ejection of their stellar envelope prior to the SN event. Alternatively, the SN ejecta themselves could have stripped the companion of its envelope (Marietta, Burrows & Fryxell 2000). As stated by Borkowski et al. (2006), at late stages of SNR evolution, the Fe-contaminated envelope would be evident as dense Fe-rich gas close to the centre of the SNR. Ultimately, this is simply a plausible scenario to explain the observed X-ray morphologies and spectra of MCSNR J0508$-6902$ and similar remnants. However, it is clear that the standard Type Ia model fails to explain the properties of these SNR s and some mechanism for a pre-explosion increase in the CSM density is essential.

Secondary evidence to support the ‘prompt’ explosion scenario can be found in the stellar content and star formation history (SFH) of the region, further supporting the scenario described above.

To understand the type of stellar environment the progenitor of this SNR came from, we make use of data from the Magellanic Cloud Photometric Survey (Zaritsky et al. 2004) to construct colour–magnitude diagrams and identify blue stars more massive than $\sim 8 \text{ M}_\odot$ within a $100 \text{ pc}$ ($\sim 6.9 \text{ arcm}$) radius of MCSNR J0508$-6902$. This allows us to see the prevalence of early-type stars close to the remnant. In the case of MCSNR J0508$-6902$, there are not many early-type stars close to the remnant (16), which are far fewer than around DEM L205 (Maggi et al. 2012), where we found 142 stars and concluded to a CC-SN origin, and more similar to SNRs such as SNR J0530$-7007$ (de Horta et al. 2012) with 13 early-type stars or J0527$-7104$ (Kavanagh et al. 2013) with 5 such stars.

Studying the SFH in the neighbourhood of MCSNR J0508$-6902$ (see Fig. 11) as derived from the map of Harris & Zaritsky (2009), we see modest levels of star formation ($\sim 10^{-3} \text{ M}_\odot \text{ yr}^{-1}$) peaking 12 Myr ago (accounting for the presence of several early-type stars as described in the previous paragraph), 100 Myr ago, and extending back from 250 Myr to 1 Gyr. The peak at a look-back time of 100 Myr is consistent with a relatively young progenitor able to produce dense CSM, but we cannot rule out an older progenitor. A ‘delayed’ progenitor older than $\sim 2.4 \text{ Gyr}$ (e.g. Maoz et al. 2011) is however less likely.

Borkowski et al. (2006) discussed the possibility of the iron-rich DEM L238 and DEM L249 remnants to be those of sub-luminous SN2002cx-like Type Ia SNe. Recent studies of this peculiar class of SNe (Jordan et al. 2012; Foley et al. 2013; Kromer et al. 2013) interpret these SNe as failed detonations of a white dwarf, which leads to relatively low initial kinetic energies ($\sim 0.13 \times 10^{51} \text{ erg}$) and ejecta masses ($\lesssim 0.5 \text{ M}_\odot$). Our estimates of $E_\text{kin}$ and $M_\text{ej}$ are at odds with these properties, suggesting that MCSNR J0508$-6902$ is likely not the remnant of a failed-detonation SN. Further studies of centrally iron-rich SNRs and discovery of new members of that class are highly desirable to address the questions of what are the progenitors of these objects, and if and how they differ from the remnants of ‘standard’ Type Ia SNe.

## 6 SUMMARY

We have undertaken a detailed multiwavelength study of a new LMC SNR candidate, MCSNR J0508$-6902$, using radio, X-ray and optical data (both imaging and spectroscopy) that confirm we have uncovered a new LMC SNR. With this analysis we infer that MCSNR J0508$-6902$ is an evolved remnant, with a large size of $\sim 74 \times 57 \text{ pc}$. We estimate the magnetic field strength to be $\sim 28 \mu \text{ G}$ and an age of $\sim 20$–$25 \text{ kyr}$. Based on the determined O/Fe ratio of the interior gas and the mass of Fe ejecta, we conclude that this remnant is of a thermonuclear SN. It is most likely that this remnant was the result of a ‘prompt’ Type Ia event, sharing similar properties to various other prompt Ia LMC SNRs.

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## REFERENCES

APPENDIX A: X-RAY BACKGROUND

In the case of single- and double-pixel events, the X-ray background comprises particle-induced and astrophysical components, as well as electronic read-out noise. The particle-induced background of the EPIC consists of three components. The first of these is the quiescent particle background (QPB), a continuum component produced by the direct interaction of penetrating high-energy particles with the detector. Depending on the energy range considered, this can be modelled by a power law not folded through the instrumental response (see Freyberg et al. 2004; Nevalainen, Markevitch & Lumb 2005; Zhang et al. 2009, for example). The second component consists of the instrumental fluorescence lines, which are due to the interaction of high-energy particles with the material surrounding the detector (Lumb et al. 2002; Read & Ponman 2003). Below 5 keV, only the Al Kα instrumental fluorescence line at 1.49 keV contributes to the EPIC-pn particle-induced background. Finally, we have residual SP contamination, which is due to SPs travelling through the telescope system to the detector. This component is quite variable with periods of high SP background filtered out of an observation event list during a standard reduction. However, there will always remain some residual SP contamination which escapes the filtering criteria and must be considered. The spectrum of these residual SPs can be described by a power law not convolved with the instrumental response (Kuntz & Snowden 2008), and can vary in both magnitude and slope. The electronic read-out noise is very significant at the lowest energies, increasing dramatically below ~0.7 keV. Since we limit our spectral analysis to the 0.4–5 keV energy range, we avoid the most affected energy range; however, some increase in the background is expected just above 0.4 keV.

A1 QPB, instrumental fluorescence lines and electronic read-out noise

To model the QPB, instrumental fluorescence lines and the electronic read-out noise, we make use of the XMM–Newton FWC data, which were collected with the detector shielded from the astrophysical and SP backgrounds. A full discussion on the FWC data of the EPIC-pn is given in Freyberg et al. (2004). The FWC data were screened using the same event filtering criteria as for the observational data, and were also vignetting corrected using the SAS task evgweight. Although the QPB is not subject to vignetting, it is necessary to apply the same correction as to the observational data. For the purposes of spectral fitting, ancillary response files are generated for a flat detector map as the effective area correction has already been applied to the event list.

To characterize the FWC spectra, and to assess the likely contribution of the QPB to our observational data, we extract spectra from the FWC data from the same regions of the detector as our SNR source and background spectra. However, these spectra must be scaled to account for differences in the integrated QPB count rates of the observational and FWC data. To perform this scaling, we considered the 10–13 keV count rates for the entire FOV of both data sets and determined a scaling factor. Unless a very bright and hard source is present in the observational FOV, the QPB continuum should dominate both data sets at these high energies. Scaling the FWC data using lower energy range count rates is not safe as there is potentially an astrophysical contribution in the observational data. We selected the energy range 10–13 keV for the scaling as Katayama et al. (2004) showed a good correlation between the 2–7 keV count rate and the 10–13 keV count rate in the FWC observations. Thus, the 10–13 keV count rates can be used to reliably scale the lower energy QPB continuum. The scaled FWC spectra are shown in Fig. A1. As previously mentioned, the background spectra were extracted from regions further off-axis than those of the SNR. The fact that these were obtained from a vignetting-corrected FWC data set has important effects on the normalizations of the QPB components. The vignetting-correction process operates on every event in the data, regardless of photon or particle origin, according to its position on the detector to account for the variation in effective area of the instrument. However, this correction is photon specific and the particle background is not subject to the same effects. This is clearly evident from Fig. A1, with the background spectra having a larger normalization than their corresponding source spectra due to the higher vignetting correction applied to events further off-axis.

We chose not to subtract these FWC spectra from the observational spectra because, as pointed out by Kuntz & Snowden (2008) and Snowden et al. (2008), the fluorescence lines of the EPIC instruments are sufficiently strong that small gain variations and...
differences in environment between FWC and filter wheel open positions can produce significant differences in the lines and correspondingly large residuals after subtraction. Hence, we chose instead to explicitly model the background in our spectral fits, using the FWC spectra shown in Fig. A1 as a basis. We fit models to the FWC spectra simultaneously, allowing only normalizations to vary, and found that the QPB continuum and electronic read-out noise components in the 0.4–4 keV range were best fitted using a broken power law (bplpower in XSPEC) not folded through the instrument response, with the energy break ($E_{\text{break}} = 0.62$ keV, the photon index for $E < E_{\text{break}}$ ($\Gamma_1$) = 1.91, the photon index for $E > E_{\text{break}}$ ($\Gamma_2$) = 0.45 and reduced $\chi^2 = 1.02$ for 1363 d.o.f. Here, $\Gamma_1$ represents the electronic read-out noise and $\Gamma_2$ the QPB. While at lower energies the electronic read-out noise component requires a more complicated model, in the 0.4–5 keV energy range we consider, a power law is sufficient. A Gaussian line was also included to account for the Al Kα instrumental fluorescence line at 1.49 keV. The fits to the FWC spectra are indicated in Fig. A1 as the solid lines. Above 4–5 keV, the defined model no longer applies as the QPB ‘flattens’. However, this is unimportant in our analysis as we can easily constrain the QPB model given that this component dominates the emission above 1.5–2 keV. In our spectral analysis of MCSNR J0508–6902, the ratios of the normalizations of the source and background QPB components were fixed to the values determined from the QPB spectra.

### A2 SP contamination

To account for the residual SP contamination in our spectra, we included the appropriate power-law model not convolved with the instrument response. However, the fits would consistently return a zero normalization for the SP component, suggesting that the SP contamination is not significant. This could be expected given the fact that there were no SP flares during the observation. However, we note that a lack of SP flares does not necessarily mean no significant SP contamination. Thus, we ran the diagnostic algorithm of De Luca & Molendi (2004)$^8$ which suggested little or no contamination by SPs. A further indication that SP contamination is minimal is that the shell spectrum above 2 keV, where we expect the SP contamination to be most noticeable, can be accounted for by the QPB component, with the normalization of the QPB components in all our fits agreeing with those of the FWC spectra within their 90 per cent confidence intervals. For these reasons, we conclude that residual SPs have no significant effect on our spectra and we ignore their contribution in our analysis.

### A3 Astrophysical background

The astrophysical background can usually be modelled with four or fewer components (Snowden et al. 2008; Kuntz & Snowden 2010). These are the unabsorbed thermal emission from the local hot bubble (LHB, $kT \sim 0.1$ keV), absorbed cool ($kT \sim 0.25$ keV) thermal emission from the Galactic halo, and an absorbed power law representing unresolved background cosmological sources ($\Gamma \sim 1.46$; Chen, Fabian & Gendreau 1997). The cool Galactic halo component likely does not contribute to the astrophysical background in this case as it is absorbed by the foreground Galactic absorbing column and, thus, we do not consider it in our background models. During our analysis, we allowed the LHB and warm Galactic halo temperatures to vary, with the spectral index of the unresolved background power law fixed at 1.46. Additionally, it was assumed that the temperature of the thermal components and the surface brightness of the thermal and non-thermal components did not vary significantly between the source and background regions. Thus, the appropriate temperature and normalization parameters were linked. To model the absorption of the Galactic halo and cosmological background sources, we used a photoelectric absorption model in XSPEC. The Galactic halo emission is subject to foreground Galactic absorption ($N_{H\text{Gal}}$), which was fixed at $7 \times 10^{20}$ cm$^{-2}$ based on the $\text{H}_1$ maps of Dickey & Lockman (1990). The unresolved cosmological background component is not only subject to this Galactic absorption, but also absorption by material in the LMC ($N_{H\text{LMC}}$). This LMC absorption component was incorporated into the fits using a vphabs model with abundance at 0.5 solar as appropriate for the LMC (Russell & Dopita 1992). However, the formal best-fitting value of $N_{H\text{LMC}}$ invariably tended to zero and was ultimately fixed as so. The astrophysical background was sufficiently accounted for by our defined model, the best-fitting parameters of which were consistent in all of our fits. The temperature of the LHB and warm Galactic halo components were (0.08–0.11) and (0.28–0.33) keV, respectively, which are compatible the expected values from Snowden et al. (2008) and Kuntz & Snowden (2010).

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$^8$ See [http://xmm2.esac.esa.int/external/xmm_sw_cal/background/epic_scripts.shtml](http://xmm2.esac.esa.int/external/xmm_sw_cal/background/epic_scripts.shtml)

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